

# Planetesimal Sizes and Mars Formation in the Magnetized Solar Nebula

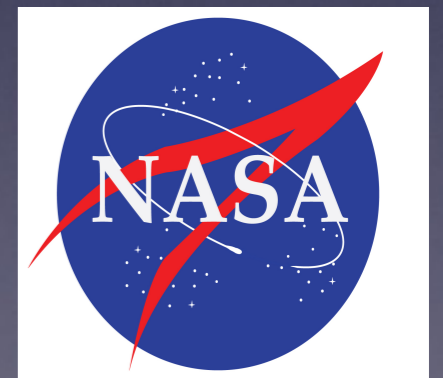
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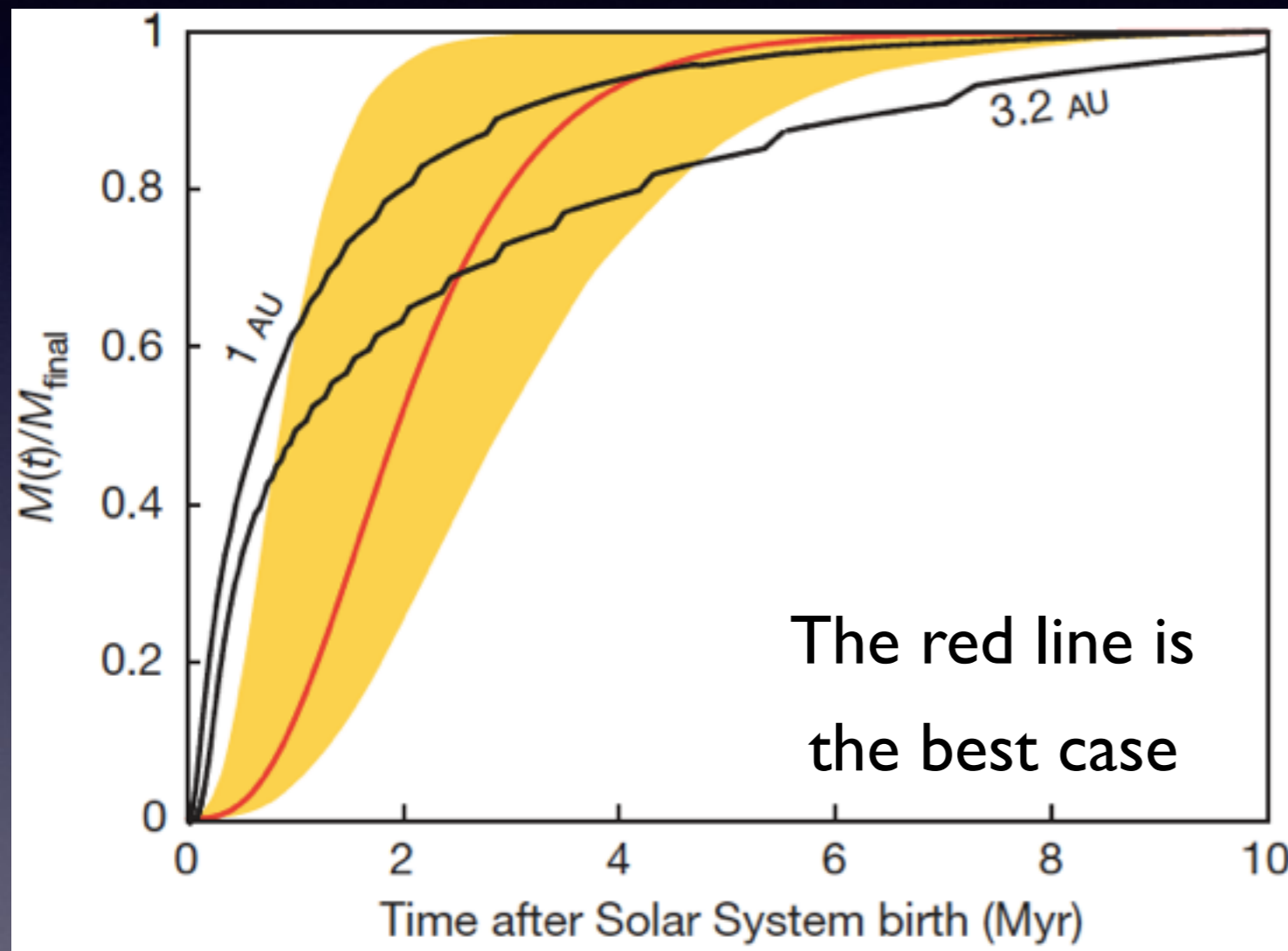
in collaboration with  
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Takaya Nozawa (NAOJ)  
Shigeru Wakita (NAOJ)



# Hf-W Chronology for Mars Formation

Dauphas & Pourmand 2011

$^{182}\text{Hf} \longrightarrow ^{182}\text{W}$  with the half life of 9 Myr



The excess of  $^{182}\text{W}$  at Mars mantle is estimated by Martian meteorites

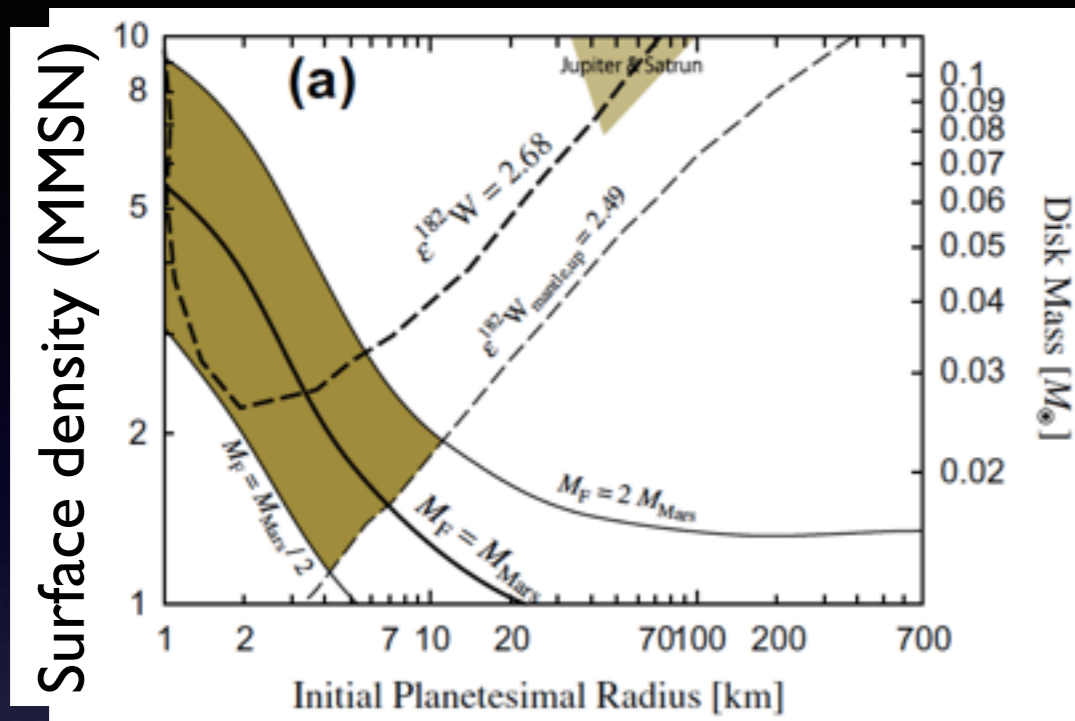
The optimal case is realized when Mars formation is completed at 2-4 Myrs after CAI formation

$^{60}\text{Fe} - ^{60}\text{Ni}$  Chronology also confirmed the short formation timescale

Tang & Dauphas 2014

# 1. Mars formed quickly after CAI formation

e.g., Kobayashi & Dauphas 2013, Morishima et al 2013, Levison et al 2015

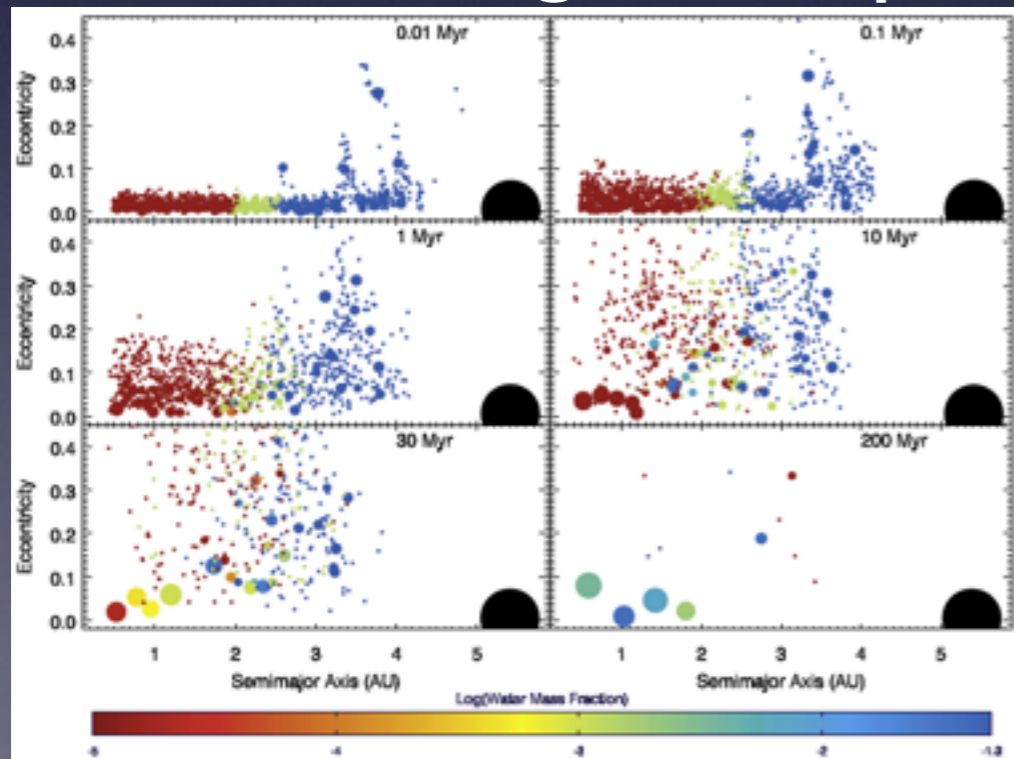


Small planetesimals  
(even pebbles)

High surface density  
: destruction/radial drift of  
such planetesimals

# 2. Mars avoided giant impacts after the nebula was gone

e.g., Raymond et al 2009, Hansen 2009, Walsh 2011



Jupiter and Saturn  
: remove planet-forming  
materials

# 3. The turbulent solar nebula was present at Mars formation



# Our Model

## Initial Conditions

- : surface density
- : magnetic fields
- : planetesimal size

## Comparison with Hf-W chronology

- : Formation history
- : excess of W



## Runaway growth

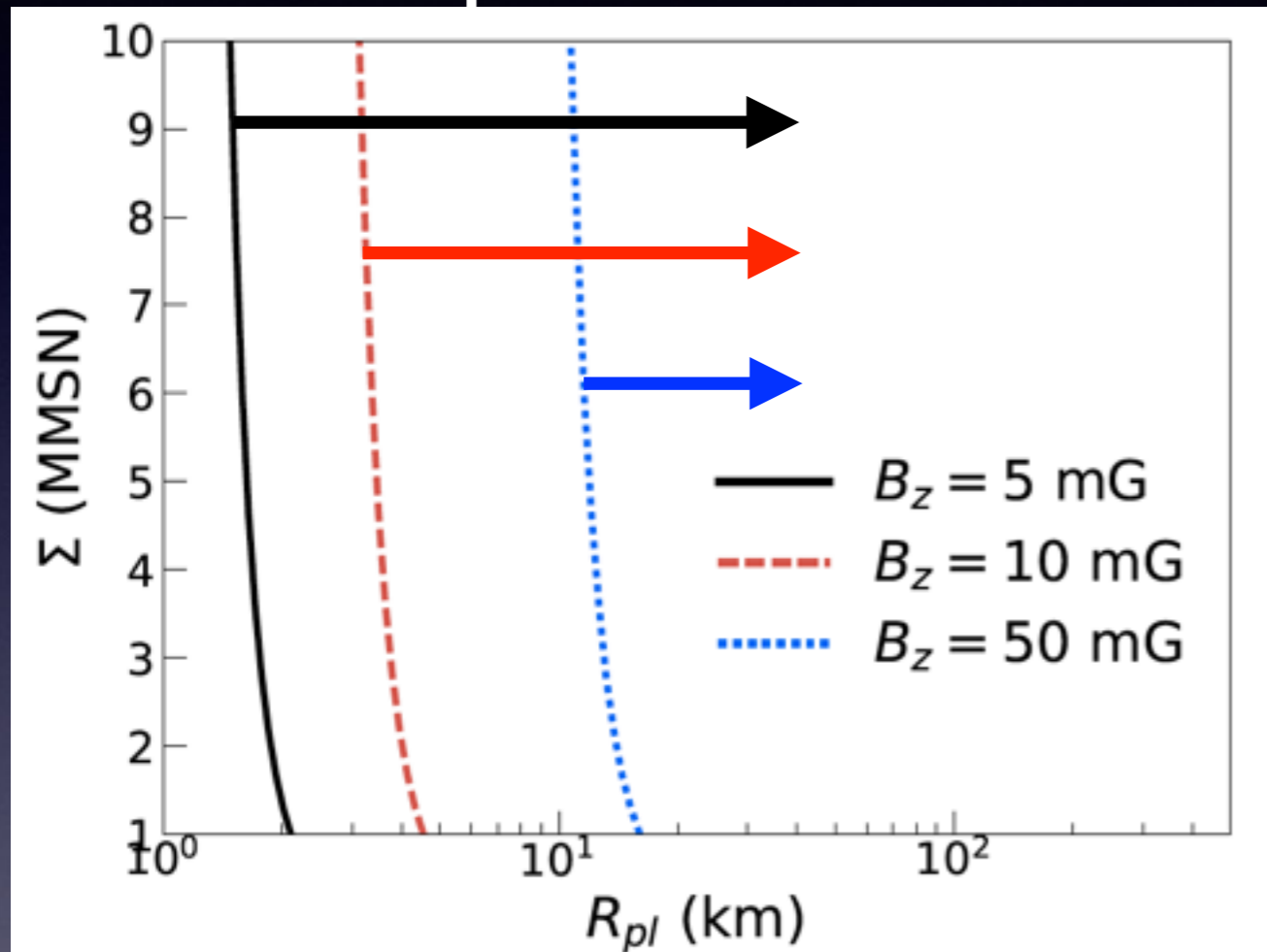
- : the minimum size of planetesimals that survive destruction (e.g., Ormel & Okuzumi 2013)

## Oligarchic growth

- : the maximum size of planetesimals that meet the short formation timescale (e.g., Chambers 2006)

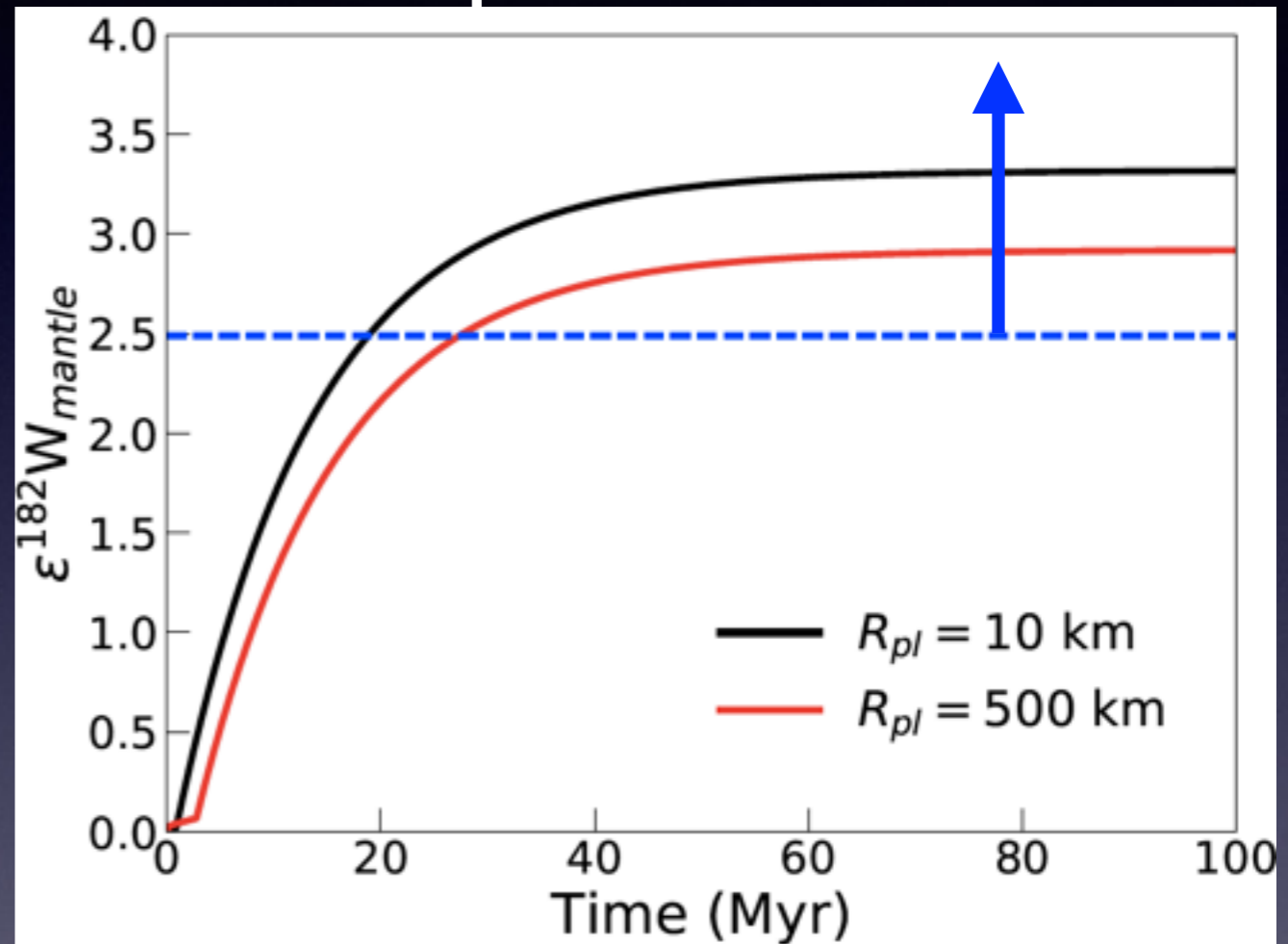
# Preliminary results

Minimum size  
of planetesimals



Larger planetesimals  
for higher B-fields

Maximum size  
of planetesimals



A wide range of planetesimals  
for a given surface density

# Next steps & Summary

- Hf-W Chronology suggests that Mars formed quickly after CAI formation
- Take into account the effect of the nebular turbulence
- develop a semi-analytical model in which the optimal values of planetesimal size and the nebular mass are specified
- will cover the larger parameter space to find out the nebular mass
- will compare other scenarios such as the narrow ring and pebble accretion