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A description of the DSN VLBI data set and of most aspects of the data analysis can be found in IERS Technical Note 17, pp. R-19 to R-32 (see also IERS Technical Note 19, pp. R-21 to R-27). The main changes in this year's analysis from last year's are simply due to including another year's data.

This year we have removed one small discrepancy between the IERS standards and our software by computing the equation of equinoxes using the mean of date obliquity rather than the true of date obliquity (see IERS Technical Note 13, pp. 30-31). We still compute the pole tide using the full value of the polar motion in the conventional terrestrial system with no "mean" value removed, since the concept of a "mean" here seems poorly defined (see IERS Technical Note 13, p. 59).

Some changes in processing strategy were tested but were not adopted for the final solution because they did not seem to significantly improve the results. This year these included (a) estimating permanent troposphere gradients at each complex, and (b) adjusting the observable uncertainty based on the scan duration (intended primarily to better account for errors in the delay rate observables induced by tropospheric variations).

Our approach to modeling the tropospheric effects on the VLBI observables was as follows. A priori dry zenith tropospheric delays were determined from barometric pressure measurements at the DSN sites, corrected for height differences between the pressure sensor and the antennas. A priori wet zenith tropospheric delays were derived from tables of monthly average wet zenith delays for each station, which are based on historical radiosonde data. The Lanyi function was used for mapping zenith tropospheric delays to observed elevations. The temperature at the top of the boundary layer, a parameter in the Lanyi function, was taken to be the 24-hour average of the surface temperature at the station. Adjustments to the wet troposphere zenith delays were estimated every two to three hours.

During calendar year 1995, the TEMPO project produced earth rotation measurements from 85 dual frequency observing sessions, with a median standard error along the minor axis of the error ellipse of 0.3 milliarcseconds (mas), and along the major axis of 1.5 mas. During 1995 the median turnaround time for TEMPO measurements, from observation to availability of earth orientation parameters, was 43 hours.

In the Tidal ERP table below, the argument conventions are those of Sovers et al. (1993). The formal errors range from 10 to 43 microarcseconds but realistic uncertainties are probably about 70 microarcseconds (one standard deviation).

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Short Period Tidal ERP Variations

Term	Period (hours)	UT1 (microseconds) Cosine	Sine	Polar Motion Amplitude (microarcseconds) prograde	retrograde	Phase (degrees) prograde	retrograde
K2	11.96724	2.2	3.2	45	76	54	221
S2	12.00000	0.8	9.8	4	120	134	312
M2	12.42060	9.8	16.8	71	251	116	277
N2	12..65835	- 1.2	1.6	22	38	103	243
K1	23.93447	11.9	22.4	169	0	139	*
P1	24.06589	- 3.1	- 3.3	77	0	316	*
O1	25.81934	-13.5	-14.6	141	0	313	*
Q1	26.86836	3.5	- 0.5	40	0	311	*

Celestial Ephemeris Pole Motion Model
(nutations relative to ZMOA-1990-2)

IAU Index	Period days	I'base	Component	Adjustment mas	Formal Error mas	Generalized Error mas
precession			Longitude	-3.00/yr	0.04/yr	0.05/yr
obliquity rate			Obliquity	-0.26/yr	0.02/yr	0.02/yr
Y-offset.			L sin eps	-1-1.29	0.21	0.24
x--offset			Obliquity	+ 5.48	0.24	0.25
1	-6798.38	In	Longitude	- 0.10	0.24	0.29
			Obliquity	- 0.06	0.07	0.07
		Out	Longitude	+ 0.21	0.15	0.18
			Obliquity	- 0.03	0.11	0.11
2	-3399.19	In	Obliquity	- 0.22	0.04	0.04
		Out	Longitude	-- 0.21	0.10	0.11
			Obliquity	+ 0.10	0.05	0.06
10	365.26	In	Longitude	- 0.22	0.05	0.05
			Obliquity	+ 0.06	0.02	0.02
		Out,	Longitude	+ 0.35	0.05	0.06
			Obliquity	- 0.02	0.02	0.0.2
9	182.62	In	Longitude	- 0.10	0.04	0.05
			Obliquity	+ 0.00	0.02	0.02
		out	Longitude	+ 0.25	0.05	0.06
			Obliquity	+ 0.06	0.02	0.02
31	13.66	In	Longitude	- 0.25	0.04	0.08
			Obliquity	+ 0.10	0.02	0.03
		Out	Longitude	+ 0.48	0.05	0.09
			Obliquity	+ 0.10	0.02	0.03
	-429.8	In	Longitude	- 0.18	0.06	0.06
			Obliquity	+ 0.03	0.02	0.02
		Out	Longitude	- 0.33	0.04	0.05
			Obliquity	- 0.19	0.02	0.02

Technical description of solution JPL 96 R 01

1 - Technique : VLBI

2 - Analysis Center : JPL

3 - Software used: MODEST

4 - Data span : Oct 78 Feb 96

5 - Celestial Reference Frame: RSC (JPL 96 R 01

a - Nature: extragalactic

b - Definition of the orientation:
The Right Ascension and Declination of OJ 287 (0851+202) and the Declination of CTD 20 (0234+285) were held fixed at the values specified in RSC (IERS) 94 C 01 .

6 - Terrestrial Reference Frame : SSC(JPL) 96 R 01

a - Relativity scale: LET (TDT=geocentric with 1 AT) The relativity model used is essentially equivalent to the " consensus model " described by Eubanks .

b - Velocity of light: 299 792 458 In/s

c - Geogravitational constant : 3.9860 0448 *10**14 m**3*s**-2

d - Permanent tidal correction : Yes

e - Definition of the origin, and

f - Definition of the orientation:
Six constraints were applied to the nine coordinates (at epoch 1993.0) of DSS 15, DSS 45, and DSS 65, such that if a seven parameter transformation (3 translations, 3 rotations, 1 scale) between the JPL 1996-1 and ITRF-93 systems were estimated by unweighted least squares applied to the coordinates of DSS 15, 45, and 65, then the resulting 3 translation and 3 rotation parts of the transformation would be zero while the scale could be nonzero and unknown in advance of computing the catalog. (When expressed as the dot product of a nine dimensional unit vector with the nine station coordinates, each constraint is assigned an a priori standard deviation of 5 mm; this does not affect the resulting coordinates but does affect the calculated formal errors, giving them a more spherical distribution than would result if either very large or very small a priori standard deviations were used.)

g - Reference epoch: 1993.0

h - Tectonic p] atemodel: 1 TRF - 93 plus adjustments

i - Constraint for time evolution:
Three-dimensional site velocities were estimated for each of the three DSN complexes. A]l stations in each DSN complex were assumed to have the same site velocity. The velocities were constrained so as to produce no net translation rate and no net rotation rate, for the net-work composed of the three DSN complexes, relative to the net motion of this network of three sites as expressed in the ITRF-93 velocity field. (When expressed as the dot product of a nine dimensional unit vector with the nine site velocity components, each constraint]s assigned an a priori standard deviation of 1.0 mm/yr; this does not affect the resulting velocity components but does affect the calculated formal errors, giving them a more spherical distribution than would result if either very large! or very smal] a priori standard deviations were used.)

'/ Earth Orientation: EOP(JPL) 96 R 01

a - A priori precession model: IAU(1976) plus adjustments

b - A priori nutation model: ZMOA - 1990-2 plus adjustments

c - Short-period tidal variations in x, y, UT1:
As part of the JPL 1996-1 cat-slog solution we estimated coefficients of a model of ERP variations at four nearly-diurnal and four nearly-semidiurnal tidal frequencies (Near ly-diurnal polar motion variations were constrained t-o have no retrograde part., thus allowing simultaneous estimation of nutations.) The reported earth rotation parameters have had these tidal frequency variations removed according to the parametric model estimated in the catalog solution. (In other words, these effects have NOT been added back in producing EOP(JPL)96 R 01.)

8 - Estimated Parameters:

a - Celestial Frame: right ascension, declination (all sources, but see 5b)

b - Terrestrial Frame: Xo, Yo, Z0, (by station) $\dot{X}$, $\dot{Y}$, $\dot{Z}$ (by site)

c - Earth Orientation: UT0-UTC and Variation of Latitude of the baseline vector
precession constant, obliquity rate, celestial pole offsets at J2000
coefficients of 23 nutation terms
coefficients of 40 diurnal and semidiurnal tidal terms in ERP

d - Others: wet zenith tropospheric delays
station clock offsets, rates, and frequency offsets

Appendix 1: Summary of TEMPO Report to IERS:

JPL. NASA's Deep Space Network operates radio telescopes in three complexes: in Australia, Spain, and the USA (California). VLBI data collected from these sites by JPL between 1978 and 1996 were analyzed for celestial and terrestrial frames and earth rotation parameters, and reported as JPL 96 R 01. The celestial frame gives coordinates for 286 radio sources and is tied to RSC (IERS) 94 C 01 through three coordinates of two sources. The terrestrial frame gives station coordinates and velocities for 10 stations in 3 sites, and is tied to ITRF-93 in both location and velocity using one station in each site. The analysis gives a time series EOP(JPL) 96 R 01 containing the UT0-UTC and Variation of Latitude of a baseline vector at a frequency of two measurements per week. Additional earth rotation information is provided in estimated corrections to precession, obliquity rate, celestial pole offsets at epoch, 23 coefficients of nutation terms, and 40 coefficients of a parametric model for the nearly-diurnal and nearly-semidiurnal tidal frequency variations of UT1 and polar motion.

Appendix 2: Operational Characteristics of TEMPO VLBI Data:

NASA's Deep Space Network (DSN) operates radio telescopes for the primary purpose of communicating with interplanetary spacecraft. The DSN has three complexes: in California, in Spain, and in Australia. The Time and Earth Motion Precision Observations (TEMPO) project uses the DSN telescopes to make rapid turnaround VLBI measurements of station clock synchronization and earth orientation in support of spacecraft navigation, which needs extremely timely, moderate accuracy earth rotation information. In TEMPO observations the raw bit streams recorded at the telescopes are telemetered to JPL for correlation, so that no physical transportation of magnetic tapes is involved. TEMPO uses the JPL-developed Block 1 VLBI system, which has a 500,000 bits/second sampling rate, with time-division multiplexing of channels. This sampling rate permits the telemetry, and thus makes rapid turnaround possible. The reduced sensitivity caused by the relatively low sampling rate in comparison to other present-day VLBI systems is largely compensated by the very large antennas and very low system noise levels of the DSN telescopes. At present the DSN nominally schedules two TEMPO observing sessions per week, one on the Spain-California (SC) baseline, and the other on the Australia-California (AC) baseline. Each session is generally 3 hours in duration (occasionally less), and records a maximum of 20 sources.

The Earth rotation results from each TEMPO measurement session are reported by specifying the UT0 and Variation-of-Latitude (DPHI) of the baseline VECTOR for that session. Each such UT0-DPHI pair has an associated error ellipse in the UT0-DPHI plane. Each such error ellipse is completely specified by the reported standard errors and correlation coefficient between UT0 and DPHI. For single baseline VLBI measurements of ERP, such as the TEMPO measurements, this error ellipse is typically quite elongated, with a ratio of major axis to minor axis of about 4:1. Therefore, for a proper interpretation of these data, it is CRUCIAL to make full use of the reported correlation coefficient. For a single-baseline VLBI estimate of earth rotation, the orientation of the error ellipse in the UT0-DPHI plane is mostly determined by the global station geometry. The direction of the minor axis of the error ellipse in the UT0-DPHI plane as predicted by the station geometry is called the transverse rotation direction, and corresponds to the motion of the baseline in the local horizontal at each station or equivalently to a rotation about an axis through the center of the earth and the midpoint of the baseline. In addition to being relatively insensitive to random measurement errors, the transverse rotation component is also relatively free of errors introduced by tropospheric modeling errors, antenna deformations, and other sources of systematic local-vertical errors.

TEMPO VLBI measurements are intended to support near-real-time knowledge of earth orientation. As a VLBI data type, the TEMPO results provide UT1 information that is stable with respect to the celestial and terrestrial reference frames. As a result, the TEMPO data are particularly effective when combined with a high time-resolution, rapid turnaround, but not inertially stable source of UT1 information. At JPL, meteorologically measured global atmosphere angular momentum values (and forecasts) are combined with geodetic ERP data, including the TEMPO VLBI results, to provide near-real-time values and short term predictions of earth orientation (see: Freedman, A. P., Steppe, J.A., Dickey, J.O., Eubanks, T.M., and Sung, L.-Y., The Short-Term Prediction of Universal Time and Length-of-Day Using Atmospheric Angular Momentum, J. Geophys. Res., 99, 6981-6996, April 10, 1994).

The quality of real time knowledge of earth orientation is critically dependent on the timeliness of the most recent measurement, even if it has relatively large uncertainty. Therefore TEMPO results are reported even when the observing session was degraded so that the measurement uncertainty is much larger than the typical TEMPO uncertainty. Thus it is important to account for the reported uncertainty accompanying each TEMPO result. Empirical RMS residuals from a set of TEMPO data will be dominated by the small number of large-uncertainty points. Therefore RMS residuals are not a good measure of the typical accuracy of TEMPO measurements. The uncertainty scaling factors for the TEMPO data developed by Richard Gross in producing the combination-of-techniques EOP series SPACE95 were in the range 1.1 to 1.4. During calendar year 1995, the TEMPO measurements had a median standard error along the minor axis of the error ellipse of 0.3 milliarcseconds (mas), and along the major axis of 1.5 mas.

TEMPO formal uncertainties have decreased dramatically from the beginning of the **program in** 1980 to the present. Thus "average" uncertainties over the full history of the program are not representative of the uncertainties of current measurements. Similarly, typical residuals over the full history are not representative of current residuals.

Typical TEMPO results from the Australia-California (AC) baseline have an error ellipse in the AC-UT0--AC-Variation-of-Latitude plane that has its major axis nearly aligned with the AC-UT0 axis and its minor axis nearly aligned with AC-Variation-of-Latitude. Thus for AC points UT0 is essentially the weak direction and residuals of order 1.5 mas are to be expected. Most of the information content of AC points is in the Variation-of-Latitude component, so failure to use the Variation-of-Latitude amounts to throwing away most of the value of the AC points. Properly used, the AC points contribute substantially to near-real-time knowledge of Polar Motion Y, and significantly to very-near-real-time knowledge of UT1.

Typical TEMPO results from the Spain-California (SC) baseline have an error ellipse in the SC-UT0--SC-Variation-of-Latitude plane that has its major axis rotated roughly 34 degrees away from SC-Variation-of-Latitude towards negative SC-UT0. Thus the SC points have a typical UT0 uncertainty of about (1.5 mas) * sin (34 degrees) = 0.8 mas. If used without considering the correlation between UT0 and Variation of Latitude, the UT0 values will have errors of order 0.8 mas, which amounts to throwing away most of the value of the SC points. To get full value from the SC points when combining them with other EOP measurements, it is best to perform a fully multivariate combination; failing this, one should at least combine one's knowledge from non-TEMPO sources of the SC-Variation-of-Latitude with the TEMPO-reported UT0-Variation-of-Latitude pair and standard errors and correlation coefficient, to get an improved SC-UT0 before transforming it to UT1. Geometrically this amounts to intersecting the angled SC error ellipse with a "small in polar motion but large in UT1" error ellipse from other sources. Properly used, the SC points contribute substantially to near-real-time knowledge of UT1.