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\setlength{\oddsidemargin} {-1cm}
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\begin{document}
\begin{center} \large\bf
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% Title:
% ex) IAU General Assembly in Japan
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A Sub-parsec Accretion Disk in NGC 4261
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\end{center}
\begin{center} \sc
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% Authors:
% e.g. A. Author$^1$, B. Author$^2$, and C. Author$^1$
%-----
D.L. Jones$^1$ and A.E. Wehrle$^2$
%-----
\end{center}
\begin{center}
%-----
% Institute:
% e.g. $^1$Institute of A, Japan <E-mail address A> \\
% $^2$Institute of B, Japan <E-mail address B>
%-----
$^1$ Jet Propulsion Laboratory, Caltech <dj@jpl.nasa.gov> \\
$^2$ Infrared Processing and Analysis Center, JPL, Caltech
<aew@ipac.caltech.edu>
%-----
\end{center}
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% Main Text: maximum 200 words
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We observed the nuclear region of NGC 4261 (3C270) with the VLBA
to determine the morphology of the central radio source on parsec
scales. Our highest angular resolution image at 8.4 GHz shows a
very narrow gap in emission just east of the radio core (on the
counterjet side) , which we interpret as an absorption feature caused
by a small, dense inner accretion disk whose width is less than 0.1
parsec. If the inclination of this inner disk is close to that of
the much larger-scale disk imaged by HST it becomes optically thin
to 8.4 GHz radiation at a deprojected radius of about 0.8 pc. Future
VLBI observations at higher frequencies should allow us to determine
the radial electron density profile of the inner disk.
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